

# Star formation in AGNs at the hundred parsec scale using MIR high resolution images

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## ABSTRACT

It has been well established in the past decades that the central black hole masses of galaxies correlate with dynamical properties of their harbouring bulges. This notion begs the question of whether there are causal connections between the AGN and its immediate vicinity in the host galaxy. In this paper we analyse the presence of circumnuclear star formation in a sample of 15 AGN using mid-infrared observations. The data consist of a set of 11.3 $\mu$ m PAH emission and reference continuum images, taken with ground based telescopes, with sub-arcsecond resolution. By comparing our star formation estimates with AGN accretion rates, derived from X-ray luminosities, we investigate the validity of theoretical predictions for the AGN-starburst connection. Our main results are: i) circumnuclear star formation is found, at distances as low as tens of parsecs from the nucleus, in nearly half of our sample (7/15); ii) star formation luminosities are correlated with the bolometric luminosity of the AGN ( $L_{AGN}$ ) only for objects with  $L_{AGN} \geq 10^{42}$  erg s<sup>-1</sup>; iii) low luminosity AGNs ( $L_{AGN} < 10^{42}$  erg s<sup>-1</sup>) seem to have starburst luminosities far greater than their bolometric luminosities.

**Key words:** active galactic nuclei – interstellar medium – star formation.

## 1 INTRODUCTION

The properties of the bulge of galaxies correlate with the properties of their super massive black holes (SMBH, Kormendy & Richstone 1995, and references therein). Arguably the most widely known among these is the correspondence between the mass ( $M_{\bullet}$ ) of the SMBH and the velocity dispersion of stars in the bulge ( $\sigma$ ), in short the  $M_{\bullet} - \sigma$  relation (Ferrarese & Merritt 2000; Gebhardt et al. 2000; Gültekin et al. 2009; McConnell & Ma 2013). Relationships such as these raise the question of whether or not there is a causal connection between the observed properties of the active nucleus and the host galaxy.

One of the possible physical links between the accretion-powered nuclear activity and its immediate vicinity is the circumnuclear<sup>1</sup> star formation, as this implies a common mechanism for fuelling the growth of both the SMBH and the stellar bulge. Early works on

the subject have shown the ubiquitous presence of star formation in Seyfert galaxies (Rodríguez Espinosa et al. 1987; González-Delgado & Pérez 1993; Cid Fernandes et al. 2001; Kauffmann et al. 2003; Riffel et al. 2007; Davies et al. 2007). There has also been a number of articles suggesting that the feedback from accretion onto the SMBH would quench star formation by heating the available gas, deterring the gravitational collapse of molecular clouds (Silk & Rees 1998; Vollmer & Davies 2013, and references therein). On the other hand, the turbulence generated by supernova explosions could be responsible for the loss of angular momentum that ultimately leads the gas to the accretion disk (Kawakatu & Wada 2008; Hopkins & Quataert 2010; Wutschik et al. 2013).

The MIR emission from polycyclic aromatic hydrocarbon (PAH) molecules has frequently been used as a tracer of star formation (e.g. Tielens 2008; Wu et al. 2009; Gallimore et al. 2010; Diamond-Stanic & Rieke 2010; Ruschel-Dutra et al. 2014). These emission bands are produced when UV photons from young stars heat molecules to temperatures of the order of 1000K. The energy is subsequently radiated by fluorescence through the many molec-

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<sup>1</sup> Here and throughout this paper circumnuclear refers to scales of the order of hundreds of parsecs or less

ular modes of oscillation and vibration. Since it is unlikely that such large molecular species would survive the intense radiation field from the active galactic nuclei (AGN, Voit 1992; Siebenmorgen et al. 2004), the analysis of their emission is virtually free from the degeneracy considerations which are relevant for ionic lines. However, it has been argued that the  $11.3\mu\text{m}$  PAH band is an indicator of recent star formation ( $\sim 10^8$  yr Díaz-Santos et al. 2010).

This study aims at investigating the relationship between the AGN activity and the recent star formation in its vicinity. It is only natural to search for a causal relation between these phenomena at the smallest possible scales, e.g. at radii of tens of parsecs from the central engine. Such regions are often enshrouded in dusty clouds, favouring observations in wavelengths other than the optical (Nenkova et al. 2008; Ramos Almeida et al. 2011; Sales et al. 2011; González-Martín et al. 2013). At present, ground based mid infrared (MIR) observations with 10m class telescopes offer the possibility of obtaining high resolution images ( $\sim 0.4''$ ), at the same time avoiding most of the extinction from the interstellar environment. Therefore, here we analyse high spatial resolution PAH emission images from two similar ground-based instruments, namely the Very Large Telescope (VLT) Imager and Spectrometer for mid Infrared (VISIR), and the CanariCam attached to the Gran Telescopio Canarias (GTC).

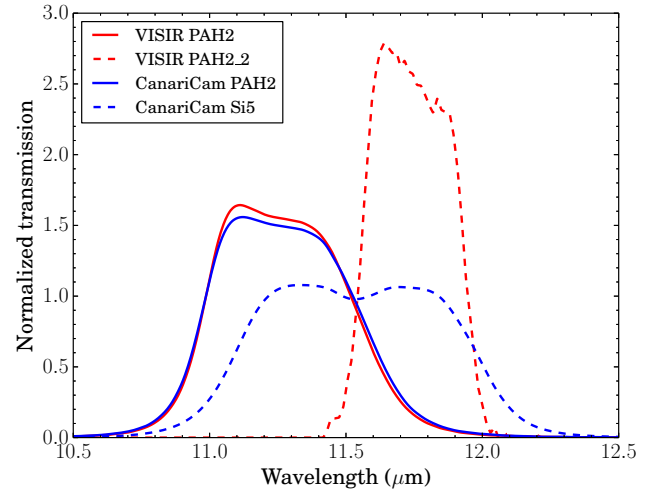
The present paper is structured as follows: in §2 we discuss the sample selection, observation strategy and reduction process; our estimates for star formation rates, as well as the morphological features of the star forming regions are examined in §3; §4 examines how the empirical data compare with numerical models for the AGN-SB relation, §5 contains a discussion of the results and finally in §6 we present our conclusions.

## 2 THE DATA

### 2.1 Sample selection

The target set for this study began with the selection of 4 galaxies, three of which were known to harbour AGNs, for observation with GTC/CanariCam, following a criteria of spectral classification diversity. The chosen sources were the Low Ionisation Nuclear Emission-Line Region (LINER) NGC 2146, the Seyfert 1 NGC 931 and the Seyfert 2's NGC 1194 and NGC 2273, the former being classified as a Compton Thick source based on X-Ray data (Guainazzi et al. 2005). The selected sources also followed a technical limit of detectability for a reasonable integration time of roughly 0.2 Jy in the  $12\mu\text{m}$  filter of IRAS. This is also the lower limit in the Extended  $12\mu\text{m}$  Galaxy Sample (Rush et al. 1993), although NGC 2273 violates the  $|b| > 25^\circ$  restriction of this catalogue. Lastly, only galaxies with projected spatial resolutions on the order of hundreds of parsecs were considered.

In order to add statistical weight to the present study, 12 sources observed with VISIR at the Very Large Telescope (VLT) were added. The data for these targets were taken from the atlas published by Asmus et al. (2014). These additional galaxies consist of all the targets that are clearly detected in both the PAH2 and Si-5 filters, and respect the



**Figure 1.** Transmission profiles of the VISIR (red) and CanariCam (blue) filters. Solid and dashed lines represent the emission and continuum filters respectively, with respect to the rest frame wavelength. The transmission integral of each filter has been normalised.

parameters of the last paragraph. The similarity between the filters in CanariCam and VISIR allows for a consistent combined analysis. The spatial scale limit applied to the proprietary data is also respected by the VISIR observations. Basic parameters for the complete sample of 15 AGNs are shown in table 1.

Our sample of nearby AGN is far from being complete both in volume and luminosity. A cross match between the AGN catalogue of Véron-Cetty & Véron (2010) and the Extended  $12\mu\text{m}$  Galaxy Sample (Rush et al. 1993) returns 44 sources with  $z < 0.03$  and spectral classification of either Seyfert or LINER. If we extrapolate on the region close to the galactic equator left out from Rush et al. (1993) there should be  $\sim 76$  galaxies that fit the description. Consequently, our sample represents close to 20% of the AGNs with  $z < 0.03$  and  $F_{12} > 0.2$  Jy.

### 2.2 Observations and data reduction

New proprietary data was acquired with CanariCam in the filters PAH2 ( $\lambda_c = 11.26\mu\text{m}$ ) and Si5 ( $\lambda_c = 11.53\mu\text{m}$ ). Additional archival data consists of VISIR observations with filters PAH2 ( $\lambda_c = 11.26\mu\text{m}$ ) and PAH2.2 ( $\lambda_c = 11.73\mu\text{m}$ ). This set of four filters, a pair for each instrument, was chosen to yield the closest possible to continuum free PAH emission images. The transmission profiles of all the filters employed in this work are shown in figure 1, where the transmission coefficient was normalised so that the total area under each curve equals unity.

The field of view for CanariCam is a rectangle measuring  $26'' \times 19''$ , with a spatial sampling of  $0.08''/\text{pixel}$ . For VISIR, however, we used data from two distinct observation modes, with square fields of view of sides  $32.5''$  and  $19.2''$ . The corresponding pixel scales are  $0.127''/\text{pixel}$  and  $0.075''/\text{pixel}$ .

The CanariCam data acquisition followed the standard recipe for MIR ground based observations, with thermal emission from the telescope and the atmosphere being re-

**Table 1.** Observation log

Target	RA (J2000)	DEC (J2000)	Sp. Type	Instrument	PAH 11.3		Si5 11.6	
					Date Obs.	$t_{\text{exp}}$ (s)	Date Obs.	$t_{\text{exp}}$ (s)
ESO 005-G004	06 <sup>h</sup> 05 <sup>m</sup> 41.7 <sup>s</sup>	-86° 37' 55.0"	Sy2	VISIR	2010-11-22	900	2010-11-22	900
ESO 138-G001	16 <sup>h</sup> 51 <sup>m</sup> 20.2 <sup>s</sup>	-59° 14' 04.2"	Sy2	VISIR	2008-03-14	600	2010-07-16	140
ESO 383-G035	13 <sup>h</sup> 35 <sup>m</sup> 53.8 <sup>s</sup>	-34° 17' 43.8"	Sy1.2	VISIR	2004-04-14	180	2010-03-10	360
IC 4329A	13 <sup>h</sup> 49 <sup>m</sup> 19.2 <sup>s</sup>	-30° 18' 33.8"	Sy1.2	VISIR	2010-03-12	180	2009-05-10	60
IC 5063	20 <sup>h</sup> 52 <sup>m</sup> 02.3 <sup>s</sup>	-57° 04' 07.6"	Sy2	VISIR	2006-05-05	180	2005-06-10	200
Mrk 1239	09 <sup>h</sup> 52 <sup>m</sup> 19.1 <sup>s</sup>	-01° 36' 43.5"	Sy1.5	VISIR	2005-01-28	1000	2006-03-12	600
NGC 253	00 <sup>h</sup> 47 <sup>m</sup> 33.1 <sup>s</sup>	-25° 17' 19.7"	Sy2/SB	VISIR	2004-12-01	1500	2004-12-01	1500
NGC 931	02 <sup>h</sup> 28 <sup>m</sup> 14.4 <sup>s</sup>	+31° 18' 41.4"	Sy1.0	CanariCam	2013-09-05	278	2013-09-05	265
NGC 1194	03 <sup>h</sup> 03 <sup>m</sup> 49.1 <sup>s</sup>	-01° 06' 13.0"	Sy2	CanariCam	2013-09-03	625	2013-09-04	199
NGC 2146	06 <sup>h</sup> 18 <sup>m</sup> 37.7 <sup>s</sup>	+78° 21' 25.3"	LINER	CanariCam	2013-09-04	139	2013-09-03	662
NGC 2273	06 <sup>h</sup> 50 <sup>m</sup> 08.6 <sup>s</sup>	+60° 50' 44.5"	Sy2ct	CanariCam	2013-09-08	625	2013-09-06	662
NGC 5128	13 <sup>h</sup> 25 <sup>m</sup> 27.6 <sup>s</sup>	-43° 01' 08.8"	Sy2	VISIR	2006-04-09	180	2006-03-15	600
NGC 5506	14 <sup>h</sup> 13 <sup>m</sup> 14.9 <sup>s</sup>	-03° 12' 27.2"	Sy1.9	VISIR	2010-02-23	180	2006-06-06	600
NGC 5995	15 <sup>h</sup> 48 <sup>m</sup> 24.9 <sup>s</sup>	-13° 45' 28.0"	Sy2/SB	VISIR	2010-07-26	1000	2010-07-26	1000
NGC 6240	16 <sup>h</sup> 52 <sup>m</sup> 58.8 <sup>s</sup>	+02° 24' 03.6"	Sy2/LINER	VISIR	2005-04-19	1800	2005-04-19	1800
NGC 7469	23 <sup>h</sup> 03 <sup>m</sup> 15.6 <sup>s</sup>	+08° 52' 25.3"	Sy1.2	VISIR	2006-07-12	180	2006-06-15	600

moved by the chopping/nodding technique. The same applies to the archival VISIR data. In the latter case, some of the nod frames which had intense detector artefacts were classified by visual inspection and discarded, resulting in a total exposure time slightly lower than the one originally intended.

The reduction process for the CanariCam data employed the RedCan pipeline (González-Martín et al. 2013), with the addition of a recently developed algorithm to register the centroid in a sequence of nod images and realign the exposures. This allowed us to obtain nearly diffraction limited images from what was otherwise a seeing limited stack. Standard stars, observed not more than an hour apart from the science exposures, were used to flux calibrate the images. The VISIR images were used as published in Asmus et al. (2014), thus we refer the reader to that paper for details about the reduction process.

### 3 IMAGING ANALYSIS

In this section we discuss the methodology and results from the analysis of the MIR images. At first the problem of isolating the location of PAH emission is investigated, followed by a description of the new observations with CanariCam. Finally photometric measurements are discussed.

#### 3.1 PAH emission maps

The simplest approach to the problem of isolating the 11.3  $\mu\text{m}$  PAH emission would be to produce maps, employing a method to compensate for the different continuum levels in both filters. The objective here is not yet to reach an accurate measurement of PAH emission, but rather to locate excesses in the PAH sampling filter with respect to the expected continuum level.

The width of MIR filters tends to be considerably larger than their counterparts in the optical range, close to half a micrometer, as can be seen in figure 1. With central wavelengths separated by distances that may be smaller than the filter's full width at half maximum (FWHM), it is clear that

there is considerable overlap between the emission and reference filters in both instruments. This is particularly true in the case of CanariCam, where the continuum reference filter has nearly 60% of its transmission curve in common with the PAH2 filter. Furthermore, the distance between the filters, although small in comparison to their FWHM, are large enough for differences in continuum levels to become noticeable. Therefore, a simple subtraction of the continuum level, as probed by the reference filter, cannot lead to an accurate estimate of the PAH emission.

Analysing tens of AGNs with already published high resolution MIR spectra (e.g. Ruschel-Dutra et al. 2014; Sales et al. 2014; González-Martín et al. 2013; Esquej et al. 2013), one finds that low resolution spectra, taken with the Infrared Spectrometer (IRS) aboard the Spitzer space telescope, show an underlying continuum that is very similar to that of the high resolution spectra. The reason for this agreement lies in the dominant role of power law and warm dust emission associated with the AGN, plus the silicate absorption bands at 9.7  $\mu\text{m}$  and 16.8  $\mu\text{m}$ . The later, even if not directly linked to the active nucleus, imprints its effect on the line of sight that leads to the AGN.

In order to produce a model of the continuum we examined the *Spitzer*/IRS spectra of our targets, available at the *Spitzer* archive<sup>2</sup>, with the spectral analysis tool PAHFIT (Smith et al. 2007). By fitting the spectrum as a combination of continuum emission, silicate absorption, and emission from molecules and ions, the code is able to return a sophisticated estimate for the continuum shape. Since we are studying galaxies, the redshift of the spectrum had to be taken into consideration when producing the model. The reader should keep in mind that *Spitzer*/IRS has a slit width of 3.6", and therefore is sampling a much larger region of the galaxy than ground based observations. Nevertheless the dominance of the nucleus over the host galaxy makes it possible to obtain a fair approximation of the continuum using low resolution spectroscopy.

We use this continuum emission as an estimate of the

<sup>2</sup> <http://sha.ipac.caltech.edu/applications/Spitzer/SHA/>

slope of the continuum under the PAH feature. Assuming that the same continuum shape holds throughout the image’s field of view, we can then produce a qualitative map of the deviation from the null hypothesis, which is having no PAH emission at all. In other words, we are building a qualitative map of PAH “excess”. The operation can be described by the equation

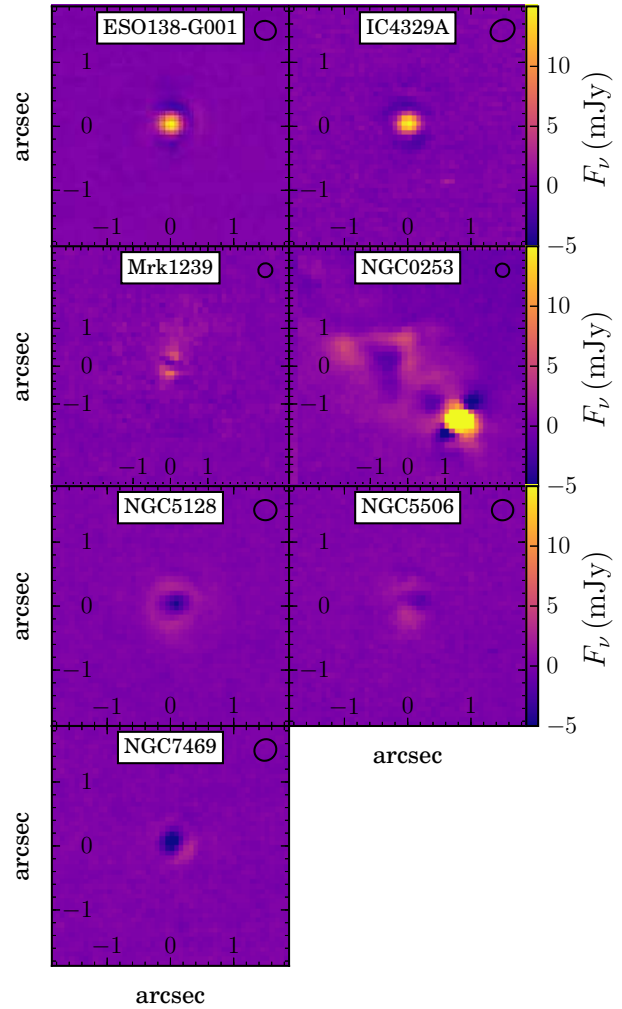
$$E = \left| I_1 - I_2 \frac{\int F_1 C d\lambda}{\int F_2 C d\lambda} \right| \quad (1)$$

where  $E$  is the difference from the null hypothesis,  $I_n$  is the image in the  $n$  filter,  $F_n$  is the normalised transmission function of the  $n$  filter and  $C$  is the function describing the semi-empirical continuum. This model is limited by the possibility of spatial variations of the continuum function, which could also produce an “excess” as the one described above. Nevertheless, there is currently no better alternative to infer the continuum slope. We would also like to emphasise that the continuum function is independently modelled for each target according to its *Spitzer*/IRS spectrum and redshift.

Employing this method we conclude that from the 15 galaxies in our sample only four show PAH emission in the image subtraction. Of these, only NGC 253 displays an extended structure, with all the other galaxies appearing as unresolved sources. The PAH emission maps that resulted from the continuum subtraction are displayed in figure 2, for the spatially unresolved and resolved sources. Notably, as can be seen in table 1, the images on each filter have not been taken in the same observing night for all the targets. Therefore, the point spread function (PSF) of the images are not naturally matched. Perhaps the most notable case of PSF mismatch are the observations of Mrk 1239, which show lobes extending beyond the FWHM of the central source. We thus refrain from assertions on morphology features that are below the wider of the two PSFs. Considering this caveat, only NGC 253 can be safely classified as having extended PAH emission.

Apart from the four galaxies with larger fluxes on the PAH images, we find two sources, namely NGC 5128 and NGC 7469, which have fluxes in the PAH images lower than the expected from the inferred continuum slope. This effect can be due to an overestimate of this slope, or to some of the emission from the molecular band “leaking” into the reference filter. This last issue is further discussed in section 4 and in the appendix, where we present our method for dealing with the effects of redshift in the relative fluxes between the filters. Additionally, NGC 5506 shows a small, almost unresolved, PAH “excess” and a region where the continuum image is more intense than expected, leading to slightly negative values in figure 2. Since the net result is close to zero, we conservatively exclude this galaxy from the list of positive detections.

It is important to keep in mind the projected scale of these images, and consequently the maximum radius of an unresolved source. For Mrk 1239, the farthest galaxy in figure 2, the PAH emitting region is at most 400 pc away from the central engine, and for the nearest target, namely NGC 253, the same region is no further than 19 pc. In the case of NGC 253, structures seen in emission are coinci-



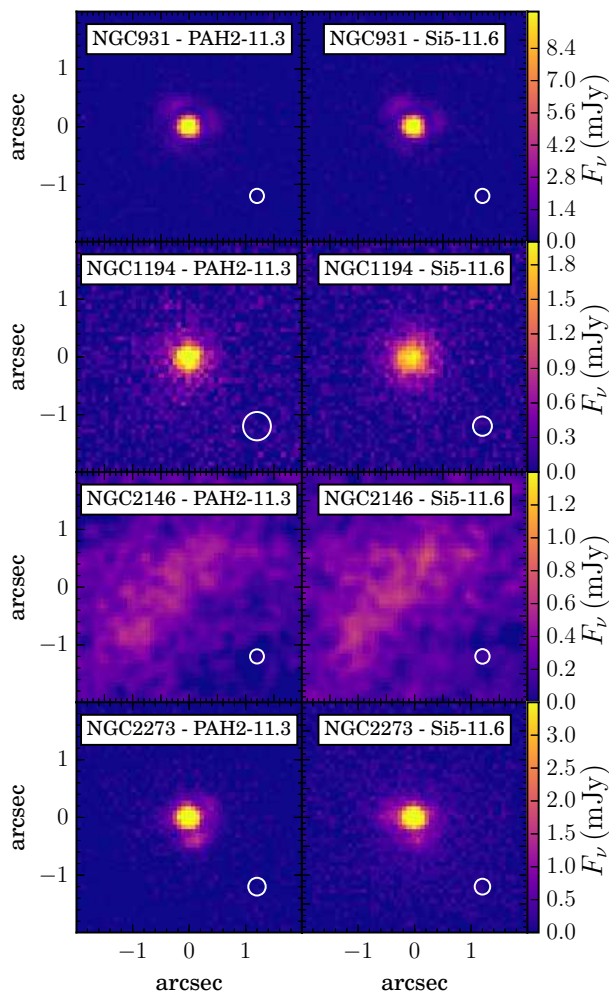
**Figure 2.** Continuum subtracted 11.3 $\mu$ m PAH images for the galaxies classified as displaying resolved or unresolved emission. The ellipses in the upper right corner of each image represent the FWHM of the standard star.

dental with H II regions already identified in the literature (Forbes et al. 2000; Lira et al. 2007).

### 3.2 Proprietary CanariCam images

Since some of our targets are having their MIR images published for the first time in this paper, we take this opportunity to examine their images in more depth. Four galaxies were observed with CanariCam: NGC 931, NGC 1194, NGC 2146 and NGC 2273. The images in the filters PAH2 - 11.3  $\mu$ m and Si5 - 11.6  $\mu$ m for the four galaxies are shown in figure 3. For display purposes, the images of NGC 2146 have been convolved with a Gaussian with  $\sigma$  of one pixel, in order to emphasise the large scale structure.

At this depth and spatial resolution, the galaxies which harbour an AGN appear as point-like sources in both filters, including the relatively close NGC 2273 ( $z = 0.006$ ). In contrast, NGC 2146 shows diffuse emission in the form of a band extending from southeast to northwest. The direction of this structure coincides with the dense dust lane identi-



**Figure 3.** Images acquired with CanariCam. With the exception of NGC 2146, for which no clear nucleus can be identified, all the other galaxies in this set appear as point like sources. In order to highlight the extended structure of NGC 2146 its image has been convoluted by a Gaussian. The faint ring like structures seen in NGC 931 and NGC 2273 are very likely instrumental artefacts (see text). In all figures North is up, East is left, and the ellipse in the lower right represents the FWHM of the standard star.

able in optical and near infrared images of this galaxy (e.g., Martini et al. 2003). Although this galaxy is classified as a LINER, no clear nucleus could be identified, and therefore it was left out of the AGN sample. We conclude that this galaxy probably has a LINER-like emission attributable to sources other than the SMBH accretion disk, such as shocks or evolved stars (Filippenko & Halpern 1984; Stasiska et al. 2008).

Only one of these four sources shows an appreciable difference in flux between filters, namely the Seyfert 2 (Sy2) NGC 1194. The similarity in filter fluxes is an expected result, since the transmission curves overlap and the distance between central wavelengths is comparable to the FWHM of the emission band we are trying to probe.

### 3.3 Artefacts in CanariCam imaging

In some of the images, particularly in the case of NGC 931 and NGC 2273, one can clearly see a pattern of three bright spots circling the central object. Since the same pattern is also visible in the standard star images they are almost certainly not real. This pattern is a combination of several effects: the hexagonal shape of GTC’s mirror and its segments; a small difference in phase between segments; and small guiding errors. The effect is not apparent in the images of NGC 1194 due to its lower signal to noise ratio.

The last of the above mentioned issues is related to the telescope guiding and the way CanariCam records the data. Each saveset is the result of nearly six seconds of chop cycles, which are stored in a buffer before being written as a file. The source position between frames varies as much as  $\sim 0.8''$  peak-to-peak, and in fact it is apparent by the structures seen in each of the savesets that such variations also occur between chop cycles. Registered stacking of the accumulated chop frames can easily solve the problem of image movement between savesets, as long as the targets are bright enough to be detected in each frame. Nevertheless, movement within each saveset requires a guiding correction frequency at least equal to the chopping frequency.

Unlike the diffraction features, which are directly related to the relative position between the detector and the primary mirror, the guiding artefacts are stochastic in nature. Thus it is impossible to eliminate them by simple comparison with the standard star. These effects, while potentially harmful to the morphological analysis, are in now way detrimental to the photometry, provided that the apertures include the stray light.

### 3.4 Photometry

Fluxes were obtained from the images through aperture photometry, considering virtual pupils with radii equal to the FWHM of the standard star employed in the flux calibration. The background levels were evaluated from an annulus with width equal to the aperture’s radius, separated from the latter by half its radius. All the sums were performed by our own routines which include treatment of partial pixels. Photometric data are displayed in table 2.

Signal to noise ratio estimates for MIR images have a few differences from their optical counterparts. The consideration that atmospheric emission follows a Poisson distribution does not hold for a chop frequency of a few tens of milliseconds. The reason is that at such short intervals there is a significant correlation between background levels in subsequent frames. Moreover, the background level is also a function of the emission from the telescope itself, making it even more time dependent. We chose to estimate the noise levels from the standard deviation of the background in the annulus.

In figures 4 and 5 we present the photometric points for the 15 galaxies in the sample, plotted along archival spectra from *Spitzer/IRS*. Since the spatial resolution of *Spitzer* is much lower than that of the data of CanariCam and VISIR, the flux in the nuclear extractions of the later is naturally smaller. We also show in these figures the continuum inferred from the spectral fitting with PAHFIT.

The photometry of all the other galaxies seems to be



**Table 2.** Photometric fluxes

ID	Target	z	$L_{2-10}$ keV (log erg s <sup>-1</sup> )	Aperture (parsec)	$f_{\text{PAH}}^1$ (mJy)	$f_{\text{REF}}^1$ (mJy)	$L_{\text{AGN}}$ (log erg s <sup>-1</sup> )	$L_{\text{PAH}}$ (10 <sup>40</sup> erg s <sup>-1</sup> )	$L_{\text{SB}}$ (log erg s <sup>-1</sup> )	SFR (M <sub>⊙</sub> yr <sup>-1</sup> )
1	ESO005-G004	0.006	41.92 <sup>a</sup>	128.7	19 ± 4	24 ± 5	42.94	13 <sup>+28</sup> <sub>-9</sub>	43.3 <sup>+0.5</sup> <sub>-0.2</sub>	0.3 <sup>+0.7</sup> <sub>-0.2</sub>
2	ESO138-G001	0.009	42.52 <sup>b</sup>	191.0	665 ± 11	640 ± 14	43.65	64 <sup>+137</sup> <sub>-43</sub>	43.9 <sup>+0.5</sup> <sub>-0.5</sub>	1.6 <sup>+3.4</sup> <sub>-1.1</sub>
3	ESO383-G035	0.008	42.41 <sup>c</sup>	164.0	305 ± 8	301 ± 19	43.52	< 4	< 42.8	< 0.1
4	IC4329A	0.016	43.70 <sup>d,e</sup>	332.2	973 ± 12	905 ± 25	45.13	855 <sup>+1849</sup> <sub>-585</sub>	45.1 <sup>+0.5</sup> <sub>-0.5</sub>	21 <sup>+46</sup> <sub>-15</sub>
5	IC5063	0.011	42.94 <sup>f</sup>	234.6	616 ± 14	809 ± 17	44.17	< 19	< 43.4	< 0.5
6	Mrk1239	0.020	40.82 <sup>g</sup>	413.2	602 ± 9	522 ± 8	41.70	< 302	< 44.6	< 7.5
7	NGC0253	0.001	40.00 <sup>h</sup>	18.7	1768 ± 97	1561 ± 114	40.83	7.8 <sup>+16.8</sup> <sub>-5.3</sub>	43.0 <sup>+0.5</sup> <sub>-0.5</sub>	0.2 <sup>+0.4</sup> <sub>-0.1</sub>
8	NGC0931	0.016	43.28 <sup>i</sup>	338.5	210 ± 10	217 ± 12	44.60	< 540	< 44.9	< 13.5
9	NGC1194	0.014	42.64 <sup>j</sup>	282.4	176 ± 17	167 ± 7	43.80	< 211	< 44.5	< 5.3
10	NGC2273	0.006	42.23 <sup>k</sup>	128.7	111 ± 5	106 ± 6	43.30	25 <sup>+55</sup> <sub>-17</sub>	43.5 <sup>+0.5</sup> <sub>-0.5</sub>	0.6 <sup>+1.4</sup> <sub>-0.4</sub>
11	NGC5128	0.002	42.31 <sup>l</sup>	37.4	823 ± 17	991 ± 13	43.40	1.3 <sup>+2.8</sup> <sub>-0.9</sub>	42.3 <sup>+0.5</sup> <sub>-0.5</sub>	0.03 <sup>+0.07</sup> <sub>-0.02</sub>
12	NGC5506	0.006	43.01 <sup>c</sup>	122.5	704 ± 15	839 ± 13	44.26	13 <sup>+28</sup> <sub>-9</sub>	43.2 <sup>+0.5</sup> <sub>-0.5</sub>	0.3 <sup>+0.7</sup> <sub>-0.2</sub>
13	NGC5995	0.025	43.53 <sup>m</sup>	521.2	290 ± 3	291 ± 5	44.92	< 23	< 43.5	< 0.6
14	NGC6240	0.024	44.26 <sup>n</sup>	504.6	152 ± 15	185 ± 12	45.87	< 116	< 44.2	< 2.9
15	NGC7469	0.016	43.02 <sup>o</sup>	330.1	393 ± 11	479 ± 9	44.27	< 68	< 44.0	< 1.7

<sup>1</sup> These flux densities are the direct result of aperture photometry over the monochromatic images, without any of the procedures for continuum subtraction or redshift compensation discussed in the text. References for the X-ray luminosities: <sup>a</sup> Winter et al. (2009);

<sup>b</sup> Piconcelli et al. (2011); <sup>c</sup> Nandra et al. (2007); <sup>d</sup> Shinozaki et al. (2006); <sup>e</sup> Bianchi et al. (2009); <sup>f</sup> Marinucci et al. (2012);

<sup>g</sup> Corral et al. (2011); <sup>h</sup> Müller-Sánchez et al. (2010); <sup>i</sup> Ueda et al. (2011); <sup>j</sup> Greenhill et al. (2008); <sup>k</sup> Awaki et al. (2009); <sup>l</sup> Shu et al. (2011); <sup>m</sup> Ueda et al. (2005); <sup>n</sup> González-Martín et al. (2006); <sup>o</sup> González-Martín & Vaughan (2012); <sup>p</sup> Brightman & Nandra (2011);

in good agreement with the *Spitzer/IRS* spectra, in the sense that the photometric points are at most equal to the spectroscopic flux. Interestingly, galaxies showing prominent 10.5μm [SiV] emission are the ones that show better agreement between nuclear photometry and host galaxy spectrum, probably due to the prevalence of the central source. Whereas in galaxies with weak [SiV] emission the AGN is correspondingly less dominant in the low spatial resolution spectrum. Targets showing very strong PAH emission in the *Spitzer/IRS* spectra tend to have nuclear fluxes well below those of the host galaxy, and even below PAHFIT's estimate for the continuum.

#### 4 STAR FORMATION RATES

In order to estimate the circumnuclear star formation rates based on the MIR photometry, we developed a method based on the difference between the photometric fluxes already discussed and careful measurements of the continuum slope. We stress that given the different redshifts of the galaxies and the large variations in silicate profile around 10μm, it is impossible to determine the flux in PAH emission features based solely on the photometric fluxes. We refer the reader to the appendix A for the technical details of this analysis.

Our results, presented in table 2, show that 7 out of the 15 galaxies in the sample have PAH emission. This number is different from the one presented in section 3.1 due to the different approach. Most importantly, simulating the PAH emission band provides information on the effects of having some of the emission sampled by the reference filter. Also, while in section 3.1 the images were compared on a pixel by pixel basis, in the present analysis only the integrated photometry is considered, thus rendering PSF mismatch problems irrelevant. Therefore, not all galaxies identified in the

image subtraction resulted in positive detections of PAH and *vice-versa*.

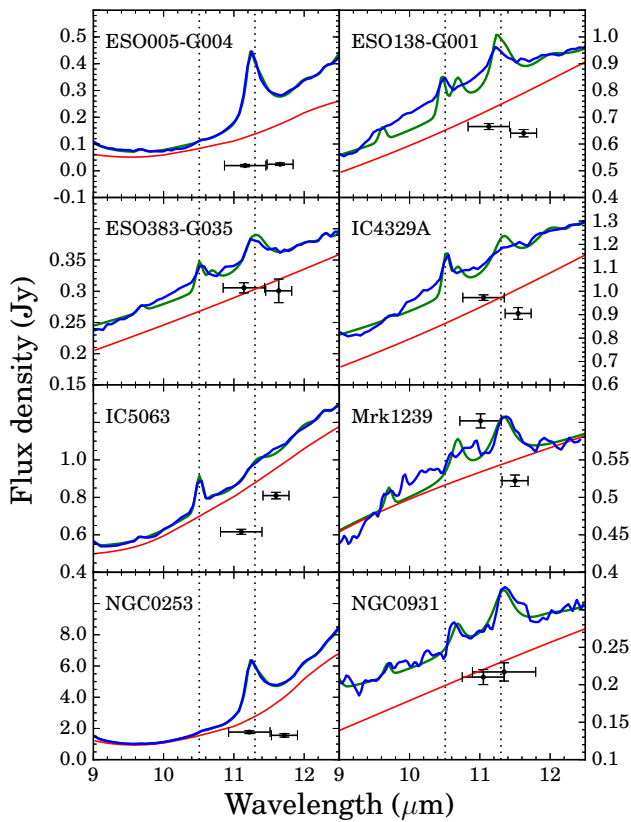
Once known, the total flux emitted in the 11.3 μm PAH flux can be used in conjunction with empirical relations for the SFR. Although this molecular feature has a less stringent correlation with SFR than the emission band at 8.6 μm (Diamond-Stanic & Rieke 2010), the features at 8.6 μm appear to be observationally suppressed in the vicinity of the AGN, while the 11.3 μm seems to be less affected (Alonso-Herrero et al. 2011). That apparent suppression should not be confused with the physical destruction of PAH molecules, but rather the dilution of PAH emission in the intense continuum emission from the AGN (Sales et al. 2010; Alonso-Herrero et al. 2014).

As previously discussed in the introduction, the 11.3 μm PAH band is a reliable MIR proxy for star formation rates. It is particularly well suited for the study of AGN due to its relative insensitivity to the radiation from the accretion disk, which is thought to destroy the molecules (Voit 1992). This is not the case of atomic fine structure lines such as the 12.8 μm [NeII], which tend to overestimate SFR in luminous AGNs (Diamond-Stanic & Rieke 2010). Other PAH features on shorter wavelengths, such as the 6.2, 7.7 and 8.6 μm features, have been shown to be suppressed in Seyferts (*ibid.*).

Based on templates of MIR spectra of starburst galaxies from Rieke et al. (2009), Diamond-Stanic & Rieke (2012) derived the relation

$$\dot{M}_* (M_{\odot} \text{ yr}^{-1}) = 9.6 \times 10^{-9} L_{\text{PAH}} (L_{\odot}) \quad (2)$$

for starbursts characterised by  $\dot{M}_* < 10 M_{\odot} \text{ yr}^{-1}$ , where  $L_{\text{PAH}}$  is the luminosity from the PAH 11.3 μm band. This equation has been evaluated for galaxies with  $10^{9.75} < L_{\text{IR}} < 10^{10.75}$ , with a dispersion of 0.28 dex. Our sam-



**Figure 4.** Spectra from *Spitzer*/IRS for the first 8 galaxies in the sample. Blue lines are the observed spectra shifted to the rest frame; red lines represent the continuum emission with the silicate absorption as considered by PAHFIT; vertical dotted lines mark the wavelength of the 10.5  $\mu\text{m}$  [S IV] line and 11.3  $\mu\text{m}$  PAH band; green lines represent the best theoretical match to the spectrum. The dots are the photometric flux points shifted to the central wavelength they are sampling at the given redshift, with error bars representing the filter's FWHM and uncertainties in the flux measurements, in the horizontal and vertical directions respectively.

ple has a median  $L_{\text{IR}} = 10^{10.65}$  with three galaxies above  $L_{\text{IR}} = 10^{11}$ , therefore we conservatively regard uncertainties in  $\dot{M}_*$  as 1 dex.

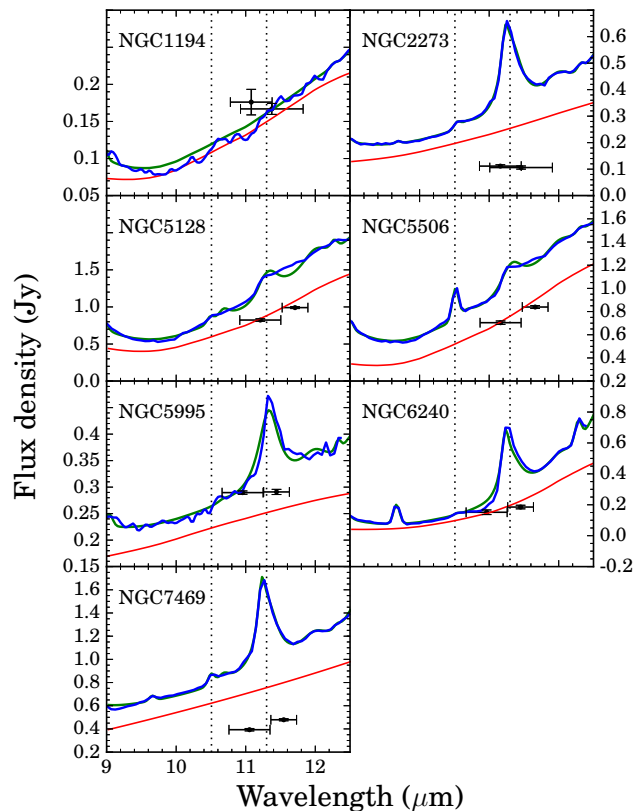
Proceeding in this line of thought, the energy-mass equivalence allows the luminosity of the circumnuclear starburst ( $L_{\text{SB}}$ ) to be expressed as function of the mass conversion rate from molecular gas to stars ( $\dot{M}_*$ ). Adding a scaling factor of 0.14 Kawakatu & Wada (2008) write  $L_{\text{SB}}$ , at a particular instant, as

$$L_{\text{SB}} = 0.14 \varepsilon \dot{M}_* c^2 \quad (3)$$

where  $c$  is the speed of light in the vacuum and  $\varepsilon$  is the mass conversion efficiency set to 0.007, which is the fraction of mass converted to energy during the fusion of Hydrogen.

These last two equations, in cgs units, result in the following relation between PAH emission at 11.3  $\mu\text{m}$  and starburst luminosities:

$$L_{\text{SB}} (\text{erg s}^{-1}) = \varepsilon c^2 2.2 \times 10^{-17} L_{\text{PAH}} (\text{erg s}^{-1}). \quad (4)$$



**Figure 5.** Same as figure 4 for galaxies 9 through 15.

Starburst luminosities, derived according to equation 4, for all the galaxies in the sample are presented in table 2.

It is important to point out that of all the seven galaxies in which we detected circumnuclear star formation, 4 have already been hinted at having such a characteristic by previous studies, namely ESO 138-G001, NGC 253, NGC 2273 and NGC 5506 (Cid Fernandes et al. 2005; Engelbracht et al. 1998; Mulchaey et al. 1996; Oliva et al. 1999). Based on the assumption that MIR observations penetrate deeper into the dusty nuclear environment, one should expect such cases. In the opposite direction, NGC 6240 and NGC 7469 have previous detections of circumnuclear star formation in the optical, but we find no evidence for it in our analysis. This could be either due to the detection limit imposed here or, in the case of NGC 7469, because the star formation structure seen by previous authors is outside the pupil used to isolate the nuclear source (see Soifer et al. 2003).

The SFR estimates presented here for the circumnuclear region of some of the galaxies are considerably smaller than previous estimates for the entire galaxy found in the literature. In the case of NGC 6240 Howell et al. (2010) report a SFR of  $148.44 M_{\odot} \text{ yr}^{-1}$  using *Spitzer*  $L_{\text{IR}}$  and GALEX FUV measurements, with the caveat that some AGN contamination may be present. The most likely explanation for the smaller value of SFR presented here is the difference in the apertures, which is almost a factor of 10 smaller when compared to that of *Spitzer*.

On the other hand NGC 253 and NGC 7469 have reported SFRs for the nuclear region with spatial resolu-

tions comparable to the ones in this paper, with values of  $2.8 \pm 0.3 M_{\odot} \text{ yr}^{-1}$  and  $2.6 - 5.1 M_{\odot} \text{ yr}^{-1}$  respectively (Ott et al. 2005; Davies et al. 2007). Our estimates for these targets are 1.1 dex and 0.6 dex smaller, therefore justifying the conservative uncertainty of 1 dex mentioned above.

## 5 DISCUSSION

A number of theoretical studies have shown that the interstellar environment in which SMBH accretion is observed also favour star formation in the circumnuclear region. For instance, Kawakatu & Wada (2008) using a semi-analytic model show that given a continuous supply of gas from outer parts of the host galaxy to the inner 100 pc, AGN luminosity will be correlated to SB luminosity while the accretion rates are high. In another semi-analytic work Neistein & Netzer (2014) show that for SB and accretion events ignited by galaxy mergers, a correlation between  $L_{\text{AGN}}$  and SFR is verified when  $L_{\text{AGN}} > 10^{42} \text{ erg s}^{-1}$ .

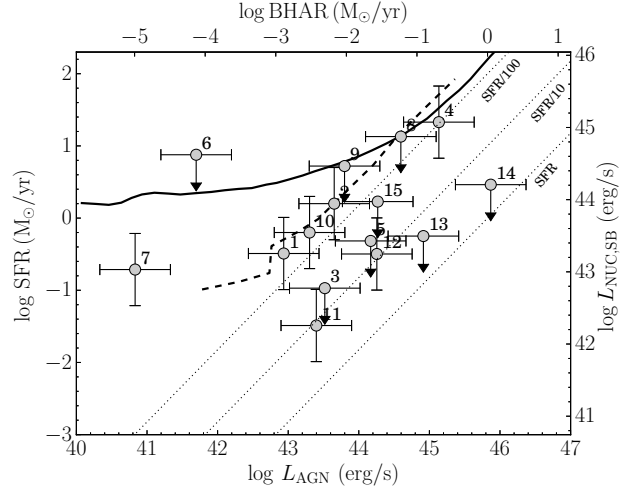
In order to investigate the possible physical connection between the AGN and the circumnuclear starburst, we compared the SFRs probed by the PAH emission to the accretion rate of the AGNs. The latter quantity was derived from the X-Ray 2-10 keV luminosity, with bolometric corrections obtained from

$$\log \left[ \frac{L_{12}}{L(2-10 \text{ keV})} \right] = \frac{1.54 + 0.24L_{12} + 0.012L_{12}^2 - 0.0015L_{12}^3}{1.54 + 0.24L_{12} + 0.012L_{12}^2 - 0.0015L_{12}^3} \quad (5)$$

(Marconi et al. 2004), where  $L_{12} = \log(L_{\text{AGN}}) - 12$  and  $L_{\text{AGN}}$  is the bolometric luminosity in units of  $L_{\odot}$ . Numerical methods were used to solve this transcendental equation, resulting in correction factors roughly between 7 and 55. Resulting bolometric luminosities for all the AGNs in our sample are also shown in table 2. We would like to emphasise that the value of  $L_{\text{AGN}}$  obtained from X-Ray radiation is only representative of the instantaneous accretion rate of the BH. On the other hand, the emission from aromatic molecules used to assess SFRs lags 150 Myr behind the main star formation event.

The points in figure 6 show the SFRs of circumnuclear regions versus the bolometric luminosities of the AGNs. In the same graph we compare our results with theoretical predictions from Neistein & Netzer (2014) by overplotting the average SFRs and  $L_{\text{AGN}}$  lines from figure 3 in the same paper. The solid line represents the average SFR for each  $L_{\text{AGN}}$  bin in Neistein & Netzer (2014) models, while the dashed line marks the average  $L_{\text{AGN}}$  for each SFR bin. We also note that the higher luminosity AGN in our sample ( $L_{\text{AGN}} > 10^{42} \text{ erg s}^{-1}$ ) are mainly located in the region these authors claim to be occupied by objects in the early stages of the starburst. Observational evidence from longer wavelengths also show a similar trend for high luminosity AGNs. For instance, Rosario et al. (2012) using 60  $\mu\text{m}$  data from the Herschel space telescope, found that local AGN luminosities correlate with SFR for  $L_{\text{AGN}} > 10^{44} \text{ erg s}^{-1}$ , thus agreeing with our findings.

Through the use of hydrodynamical simulations Hopkins & Quataert (2010) have predicted correlations between black hole accretion rates (BHAR) and SFR. They



**Figure 6.** Relation between AGN luminosity and the nuclear star formation rate. The numbers correspond to the identification on the first column of table 2. Dotted lines mark accretion rates in terms of the SFR, in ratios of BHAR = SFR, BHAR = SFR/10 and BHAR = SFR/100. The solid line and dashed lines are reproductions of the semi-analytic models in Neistein & Netzer (2014, see text).

conclude that the SFRs are close to BHARs when considering only the nuclear region ( $R < 10 \text{ pc}$ ). The spread in the correlation increases as larger radii are considered, at the same time that BHARs correspond to smaller fractions of SFRs. In our analysis we have not compensated the SFRs for the different projected areas they represent because the majority of our sources is unresolved. As a result, some of the dispersion in figure 6 can be directly linked to comparing different proportions of the host galaxy.

In figure 6 we also indicate the corresponding  $L_{\text{SB}}$  for the calculated SFRs, meaning the luminosity due exclusively to young stars. The most striking feature is that there are no galaxies where  $L_{\text{AGN}} > 100 \times L_{\text{SB}}$ . In other words, all the AGN harbouring galaxies in our sample, even the most energetic ones, have starbursts radiating at least 1% of the energy from the central engine. At the same time, the high luminosity AGNs have circumnuclear starbursts that at most match the energy output of the central source. Moreover we find that two of the seven well constrained galaxies lie within uncertainty limits of the 1-to-1 line, meaning that a significant fraction of these sources have nearly as much energy coming from the AGN as from the circumnuclear starburst. This reinforces the importance of isolating the AGN emission from star formation at the MIR.

Concerning the low luminosity AGNs, there is one galaxy which has more than one hundred times more energy coming from young stars than from the AGN. This is in agreement with recent studies showing the huge importance of the circumnuclear star forming components at all wavelengths to understand LLAGN (e.g. González-Martín et al. 2014). Conversely, we find no examples of low luminosity AGNs with  $L_{\text{SB}} \ll L_{\text{AGN}}$ . Regarding the relationship between the luminosities from the different phenomena, our data agrees with the theoretical predictions of Kawakatu & Wada (2008) and Wutschik et al. (2013), which argue that low luminosity AGNs should not display a correlation between circumnuclear star formation and AGN



activity. According to these authors, this lack of correlation represents a stage where the circumnuclear disk has become gravitationally stable, thus ceasing the accretion by the SMBH, but still has enough gas to form stars.

Diamond-Stanic & Rieke (2012) have found a positive correlation between BHAR and the SFR in radii averaging to 300 pc, throughout their sample, using the [O IV] as an indicator of the former. In that study, only two of the four points with  $L_{\text{AGN}} < 42 \text{ erg s}^{-1}$  fall within the correlation, leaving a further two points below it. This result contrasts with our analysis in the sense that we would expect to see higher circumnuclear star formation in the low luminosity regime. Using spectra rather than images, Esquej et al. (2013) arrived at circumnuclear star formation detection rates similar to ours, almost half of their sample of 23 AGNs. We have four objects in common with their work, and our estimates for 11.3  $\mu\text{m}$  PAH luminosity agree with theirs within the uncertainties.

## 6 CONCLUSIONS

We analysed MIR images of a sample of 15 AGNs in two adjacent filters (PAH2 and Si5 in the CanariCam, and PAH2 and PAH2.2 at VISIR) to study the occurrence of circumnuclear star formation via the 11.3  $\mu\text{m}$  PAH emission band. Three of these 15 galaxies were observed with GTC/CanariCam, and the remaining 12 with ESO/VISIR, with images taken from the atlas published by Asmus et al. (2014). We have also presented new high resolution MIR images for NGC 2146, which was eventually left out of the remaining analysis due to the lack of a clearly detectable nucleus. Our main results are as follows.

- Circumnuclear star formation, at distances as low as 20 pc from the nucleus, was detected in seven out of the 15 galaxies. Among these seven galaxies, four show unresolved emission, NGC 253 has clearly identifiable star formation nodes, and ESO 005-G004 and NGC 2273 have not been detected in the image subtraction.
- The luminosity from the circumnuclear starburst correlates with the bolometric luminosity of the central engine only for AGNs with luminosity higher than  $10^{42} \text{ erg s}^{-1}$ . In the lower luminosity regime, we find that the radiative energy output due to star formation tends to be higher than the central engine.

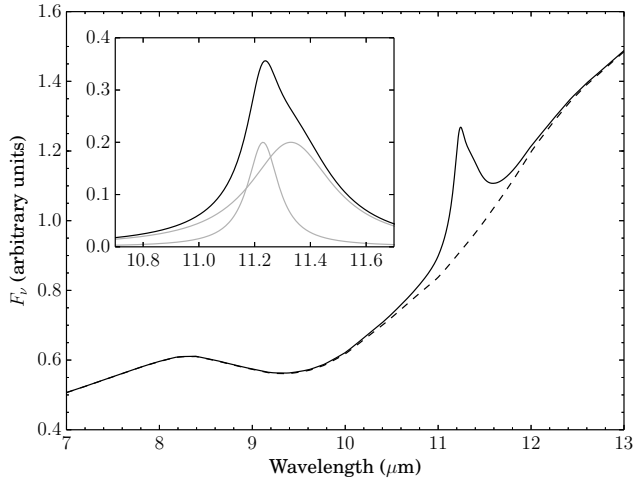
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**Figure A1.** Example of simulated spectrum used to study the effects of redshift and filter subtraction on the estimate of PAH emission. The spectrum that would be observed (*solid black line*) is the result of the sum of a continuum emission absorbed by silicate grains (*dashed line*) and the emission from PAH molecules. The smaller plot window highlights the composition of the 11.3  $\mu\text{m}$  by two distinct Drude profiles (*grey lines*.)

## APPENDIX A: ESTIMATING TOTAL INTENSITY OF THE PAH BAND

In this section we will discuss the central issue concerning the determination of SFR with the available data, which mainly consists of disentangling PAH and continuum emission with two overlapping filters. Let us assume that the specific portion of the spectrum in which we are interested, can be well represented by a continuum plus the PAH emission band

$$S(\lambda) = C(\lambda) + I(\lambda). \quad (\text{A1})$$

The later is described here by the sum of two Drude profiles parametrised by their central intensities as in Smith et al. (2007):

$$I_\nu = \frac{b_r \gamma_r}{(\lambda/\lambda_r - \lambda_r/\lambda)^2 + \gamma_r^2}. \quad (\text{A2})$$

where  $b_r$  is the central intensity,  $\gamma_r$  is the fractional full width at half maximum  $\text{FWHM}/\lambda_r$  and  $\lambda_r$  is the central wavelength (Smith et al. 2007). The values of  $\lambda_r$  and  $\gamma_r$  that define the first component are 11.22  $\mu\text{m}$  and 0.012, respectively, while 11.33  $\mu\text{m}$  and 0.032 represent the second component. Just as in section 3.1, the spectral decomposition code PAHFIT (Smith et al. 2007) was employed to obtain individual models for the continuum of each galaxy based on *Spitzer*/IRS spectra. Figure A1 shows an example of a modelled spectrum, with the average continuum of our sample. In this example, and in all subsequent analysis, the central intensity of both Drude profiles was kept equal. Tests were performed to evaluate the impact of different ratios of central intensities, and no large differences were found.

The total flux  $f_n$  measured in each filter is essentially a weighted average of the spectrum, with the weight for each wavelength given by the filter's transmission function. Thus

$$f_n = \int F_n(\lambda) S(\lambda) d\lambda. \quad (\text{A3})$$

where  $F_n$  is the normalised transmission curve. It is easy to see that, within the framework of this model, the difference in flux between filters, that is  $f_{\text{PAH}} - f_{\text{REF}}$ , can be written as function of only the central intensity of the PAH emission profiles, and the redshift of the galaxy:

$$D(b_r, z) = \int_{\lambda_0}^{\lambda_1} [C_\nu(\lambda, z) + I_\nu(b_r, \lambda, z)] \times [F_1(\lambda) - F_2(\lambda)] d\lambda \quad (\text{A4})$$

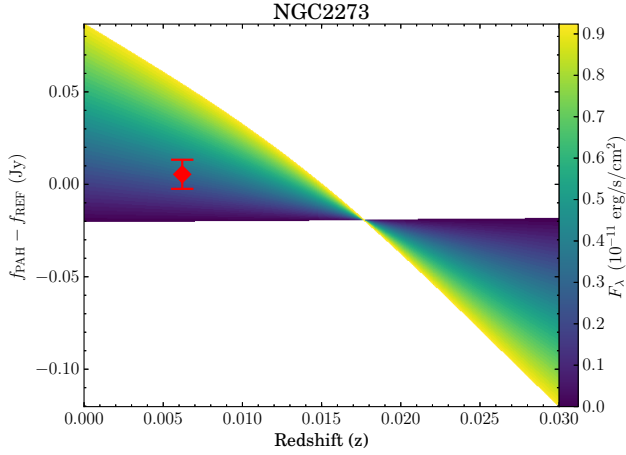
where  $C_\nu(\lambda, z)$  is the continuum emission,  $I_\nu(b_r, \lambda, z)$  is the PAH emission in the 11.3  $\mu\text{m}$  band, and  $F_1(\lambda)$  and  $F_2(\lambda)$  are the normalised transmission functions for the PAH and reference filters respectively. The total intensity of the PAH band is given by the sum of the total intensity of each component, which in turn is the integral of equation A2:

$$F_{\text{PAH}} = \frac{\pi c}{2} \left( \frac{b_1 \gamma_1}{\lambda_1} + \frac{b_2 \gamma_2}{\lambda_2} \right) \quad (\text{A5})$$

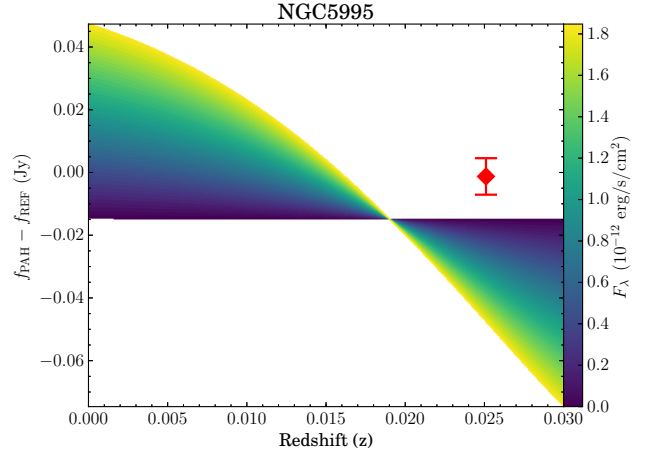
where the sub indexes represent the different components. Therefore, once equation A4 is solved for a particular central intensity  $b_r$ , and assuming the redshift  $z$  is known, we can use the result in equation A5 to calculate the total flux in the 11.3  $\mu\text{m}$  PAH band. Numerical methods were employed to investigate the solutions of A4, and figures A2 through A4 show three distinct examples of typical outcomes.

In figure A2 we show a map of solutions in the plane  $f_{\text{PAH}} - f_{\text{REF}}$ , which we assume to be equal to  $D(b_r, z)$  vs.  $z$ , for the inferred continuum of NGC 2273 and the CanariCam filters. The red diamond in this figure represents the measured filter difference, along with 1- $\sigma$  error bars. The colour scale represents the total flux of the simulated PAH band, from deep blue for no emission to bright green for the maximum simulated emission. The latter is completely arbitrary and was chosen for readability purposes. In the case of NGC 2273, the almost flat slope of the continuum has no clear impact on the difference between fluxes, for any of the sampled redshift coordinates. Nevertheless, it is already clear that even a modestly negative value of  $D$  is still compatible with a PAH emission of about 0.2  $\text{erg/s/cm}^2$ . At a redshift of 0.017 the situation is reversed, because the reference filter ( $f_{\text{REF}}$ ) begins to sample more of the PAH band than the original PAH filter. Thus, for the same continuum shape and for the same filters, the PAH emission of a galaxy with redshift  $z > 0.017$ , would cause  $f_{\text{PAH}} - f_{\text{REF}}$  to be lower than -0.2 Jy.

Figure A3 shows the example of NGC 5128, which has a notably steeper continuum (see fig. 5). At the redshift of this galaxy the flux of PAH band would have to be almost  $8 \times 10^{-10} \text{ erg/s/cm}^2$  for the difference between fluxes to be zero. The negative flux seen in the image subtraction of section 3.1 is therefore completely compatible with PAH emission, although the former alone is not enough to warrant an emission detection. The slowly rise in the minimum value of  $f_{\text{PAH}} - f_{\text{REF}}$  that can be seen in figure A3 is caused by the change in continuum slope as one moves towards longer wavelengths, leaving the silicate absorption band at 9.7  $\mu\text{m}$ .

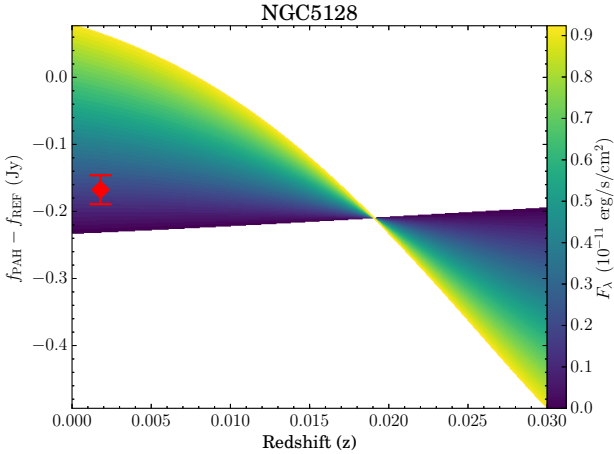


**Figure A2.** Analysis of the effects of redshift and total intensity on the difference between the filters on CanariCam. At a redshift of 0.017 the reference filter begins to have a higher flux than the emission filter even for the highest intensities allowed in the models.



**Figure A4.** Same as figure A2 for the galaxy NGC 5995

many local AGN (Sturm et al. 2005; Thompson et al. 2009; Mason et al. 2012; Ruschel-Dutra et al. 2014), and are predicted by radiative transfer models based on clumpy tori even for Seyfert 2's (Nenkova et al. 2008).



**Figure A3.** Same as figure A2 for the galaxy NGC 5128

Finally we discuss the case of the NGC 5995, the galaxy with the highest redshift in our sample. This galaxy is well beyond the threshold where the reference filter has more of the PAH band than the original PAH filter, which in the case of VISIR happens at  $z = 0.019$ . A positive detection would therefore be characterised by  $f_{\text{PAH}} - f_{\text{REF}} \leq -0.015$  Jy. However the photometric measurements show practically no difference between the fluxes in both filters, thus placing NGC 5995 in a region which would imply a negative PAH flux, or a more likely misrepresentation of the underlying continuum. We classify such cases as non-detections and report PAH fluxes and derived quantities as upper limits, based on the photometric uncertainty. As for the reasons that might have thus affected the continuum, the most probable is a simple overestimation of its declivity. Another possibility would be the presence of silicate emission at the  $9.7 \mu\text{m}$  band, rather than the more usual absorption. These emission features are a common occurrence among QSOs (e.g. Hao et al. 2005), and have also been identified in